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KiDS+VIKING-450: Improved cosmological parameter constraints from redshift calibration with self-organising maps

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ABSTRACT

We present updated cosmological constraints for the KiDS+VIKING-450 cosmic shear data set (KV450), estimated using redshift distributions and photometric samples defined using self-organising maps (SOMs). Our fiducial analysis finds marginal posterior constraints of $S_8 \equiv \sigma_8 \sqrt{\Omega_m/0.3} = 0.716^{+0.043}_{-0.038}$; smaller than, but otherwise consistent with, previous work using this data set ($|\Delta S_8| = 0.023$). We analyse additional samples and redshift distributions constructed in three ways: excluding certain spectroscopic surveys during redshift calibration, excluding lower-confidence spectroscopic redshifts in redshift calibration, and considering only photometric sources which are jointly calibrated by at least three spectroscopic surveys. In all cases, the method utilised here proves robust: we find a maximal deviation from our fiducial analysis of $|\Delta S_8| \leq 0.011$ for all samples defined and analysed using our SOM. To demonstrate the reduction in systematic biases found within our analysis, we highlight our results when performing redshift calibration without the DEEP2 spectroscopic data set. In this case we find marginal posterior constraints of $S_8 = 0.707^{+0.046}_{-0.042}$; a difference with respect to the fiducial that is both significantly smaller than, and in the opposite direction to, the equivalent shift from previous work. These results suggest that our improved cosmological parameter estimates are insensitive to pathological misrepresentation of photometric sources by the spectroscopy used for direct redshift calibration, and therefore that this systematic effect cannot be responsible for the observed difference between S_8 estimates made with KV450 and Planck CMB probes.

Key words. cosmology: observations – gravitational lensing: weak – surveys

1. Introduction

Estimation of cosmological parameters using tomographic cosmic shear requires accurate calibration of source redshift distributions. For Stage III cosmic shear surveys such as the Kilo Degree Survey (KiDS; [Kuijken et al. 2019](#)), the Dark Energy Survey (DES; [Flaugher et al. 2015](#)), and the Hyper-Suprime Camera Wide-Survey (HSC; [Aihara et al. 2018](#)), coherent biases on the order of $\Delta\langle z \rangle = \langle z \rangle_{\text{est}} - \langle z \rangle_{\text{true}} \sim 0.04$ are enough to cause significant shifts in estimated cosmological parameter estimates (see, e.g., [Hildebrandt et al. 2017](#)). Systematic shifts of this nature are important given the observed (currently mild) tension between cosmological parameters estimated using KiDS weak lensing and cosmic microwave background (CMB) studies ([Planck Collaboration et al. 2018](#)). For this reason, considerable effort has been invested in developing, testing, and optimising redshift calibration methodologies for cosmic shear. These methods can typically be grouped into three categories: those which utilise cross-correlation (see, e.g., [Schneider et al. 2006](#); [Newman 2008](#); [McQuinn & White 2013](#); [Morrison et al. 2017](#)), stacking of individual redshift probability distributions (see, e.g., [Hildebrandt et al. 2012](#); [Hoyle et al. 2018](#); [Tanaka](#)

[et al. 2018](#)), or direct calibration using spectroscopic redshift training samples (see, e.g., [Lima et al. 2008](#); [Hildebrandt et al. 2017, 2020](#); [Buchs et al. 2019](#); [Wright et al. 2020](#)). Recently, though, steps have been taken towards constructing a fourth, hybrid category which leverages both cross-correlation and direct calibration ([Rau et al. 2019](#); [Sánchez & Bernstein 2018](#); [Alarcon et al. 2019](#)).

The methodological differences, and implicit assumptions, between these estimation/calibration methods mean that they are each susceptible to subtly different biases and systematic effects. For direct calibration methods, the completeness and pre-selection of the spectroscopic training sample has been of particular concern (see, e.g., [Gruen & Brimiouille 2017](#); [Hartley et al. 2020](#)). In [Wright et al. \(2020\)](#) we developed an updated implementation of the direct calibration procedure utilising self-organising maps (SOMs; [Kohonen 1982](#)), which we found to be less susceptible to bias than previous implementations. We achieved this by the direct flagging and removal of photometric sources which are not directly associated with a spectroscopic calibrator, thereby constructing a sample of fully represented photometric sources and an associated redshift distribution: the ‘gold’ sample.

In this letter we apply the methodology of Wright et al. (2020) to the KiDS+VIKING-450 data set of Wright et al. (2019), and perform a tomographic cosmic shear analysis akin to that of Hildebrandt et al. (2020). The data set used is described in section 2, as is the definition of our various photometric and spectroscopic analysis samples. Our results are presented in section 3, and we summarise the results presented in this letter in section 4.

2. Data set and Analysis Methodology

The KiDS+VIKING-450 data set (hereafter KV450) is presented in Wright et al. (2019), and Hildebrandt et al. (2020, hereafter H20). We utilise the cosmic shear data products from H20 with `lensfit` shape measurements (Miller et al. 2007, 2013), spectroscopic training samples (Vanzella et al. 2008; Lilly et al. 2009; Popesso et al. 2009; Balestra et al. 2010; Le Fèvre et al. 2013; Newman et al. 2013; Kafle et al. 2018), and BPZ photometric redshifts from Benítez (2000), as well as the core of the H20 parameter inference pipeline; we update only the redshift distributions using the new direct redshift calibration methodology of Wright et al. (2020, hereafter W20). Our code is released as a stand-alone analysis package¹, with a wrapper pipeline² to perform the analyses presented in this work. We provide the details of these cosmological analysis pipelines in Appendix A.

In this analysis we utilise a range of differently compiled spectroscopic data sets to construct redshift distributions and photometric source ‘gold samples’ for cosmic shear analysis. A full description of the methods used to construct these redshift distributions and gold samples is presented in W20. Briefly, we utilise self-organising maps (SOMs), trained on the colours and magnitudes of sources in the various spectroscopic data sets, to associate photometric galaxies to spectroscopic galaxies with known redshift. The SOM associations are based on spectroscopic and photometric source assignments to specific SOM cells (galaxies falling within the same SOM cell are assigned to the same associations; see Appendix B of W20) and therefore links photometric and spectroscopic sources with self-similar colour and magnitude properties. Using these associations, we are able to re-weight the spectroscopic redshift distribution to approximate the (unknown) photometric galaxy redshift distributions. Simultaneously, this allows us to flag and remove photometric data which are not associated to spectra (and therefore which are not represented by the re-weighted redshift distributions).

W20 demonstrate that their redshift calibration methodology is less susceptible to systematic biases in redshift distribution reconstruction, when compared with previously incorporated methods used by KiDS (Hildebrandt et al. 2017, H20). We are able to further this analysis, also using the simulations of van den Busch et al. (2020), and estimate biases introduced by calibrating redshift distributions with different spectroscopic calibration samples. This allows us to extend our cosmological analysis to include more informative priors on the redshift distribution bias parameters, rather than incorporating only the zero-mean priors used by H20. Details of these new non-zero-mean priors ($\delta z \neq 0$) are given in Appendix B, and we adopt the use of these improved priors for the majority of our analy-

ses, where possible. In cases where we are unable to derive updated redshift priors (because of limitations in the simulations of van den Busch et al. 2020), we revert to using the broader, zero-mean priors ($\delta z = 0$) from H20.

Finally, the construction of our gold photometric source subsamples requires the simultaneous recalibration of both multiplicative and additive shear measurement bias parameters. While we are able to perform the additive shear bias measurement on-the-fly within our cosmology pipeline, computation of the multiplicative shear biases is more involved. We therefore pre-compute the required multiplicative shear bias values, using the methodology and simulations of Kannawadi et al. (2019), for each of our photometric gold samples. These bias parameters are also given in Appendix B.

2.1. Analysis samples

In this work, we perform cosmic shear parameter estimation using a number of different photometric gold samples, redshift distributions, and priors. Following the analysis of W20, our fiducial analysis defines the gold sample as being those photometric data which are associated with one or more sources within the full KV450 spectroscopic compilation, and whose spectroscopic-to-photometric associations satisfy an empirically derived quality requirement. This requirement filters out associations whose mean photometric redshift $\langle Z_B^p \rangle_i$ (from template fitting with BPZ; Benítez 2000) catastrophically disagrees with the mean spectroscopic redshift of the association $\langle z_{\text{spec}}^s \rangle_i$. Such a criteria was found in W20 to correlate strongest (in their simulations) with true bias of photometric-to-spectroscopic associations, outperforming (in particular) the measured and true spectroscopic redshift dispersions per association (which were found, counter-intuitively, to be uncorrelated with bias in the association mean redshift estimates). We therefore implement the quality requirement:

$$|\langle z_{\text{spec}}^s \rangle_i - \langle Z_B^p \rangle_i| \leq \max \left[5 \times \text{nMAD}(\langle z_{\text{spec}}^s \rangle - \langle Z_B^p \rangle), 0.4 \right], \quad (1)$$

for each of the $i \in [1, N]$ association sets, where z_{spec}^s is the spectroscopic redshift of the spectroscopic sources, Z_B^p is the photometric redshift of the spectroscopic sources, and Z_B^p is the photometric redshift of the photometric sources. This requirement is the same as presented in W20, except that we have imposed a floor on the threshold which defines catastrophic failure; we take as our threshold the maximum of 0.4 and five times the $z_{\text{spec}}^s - Z_B^p$ dispersion (determined using the normalised median absolute deviation from median; nMAD ³). Redshift distributions are then calculated per tomographic bin ($Z_B \in (0.1, 0.3], (0.3, 0.5], (0.5, 0.7], (0.7, 0.9], (0.9, 1.2]$), as are the photometric gold samples.

In addition to our fiducial analysis, we explore three gold samples constructed from spectroscopic compilations excluding the zCOSMOS, VVDS, and DEEP2 data sets, respectively. We implement these samples both to compare with similar samples run by H20, and to test the sensitivity of our results to pathologically under-representative spectroscopy. Further, we construct one gold sample (‘spec-quality4’) using only spectra which have the highest quality

³ $\sigma_{\text{nMAD}} = 1.4826 \times \text{med}(|x - \text{med}(x)|)$. The pre-factor ensures normal consistency; that is $E[\text{nMAD}(x_1, \dots, x_n)] = \sigma$ for $X \sim N(\mu, \sigma^2)$ and large n .

¹ <https://www.github.com/AngusWright/CosmoPipe>

² <https://www.github.com/AngusWright/CosmoWrapper>

flags from their various surveys (referred to as $nQ \geq 4$ spectra, which have $\geq 99.5\%$ confidence), to test the sensitivity of our analysis to spectra with a slightly higher likelihood of catastrophic failures. Finally, we construct a highly restrictive gold sample ('multispec3') which consists only of sources which reside in associations containing spectroscopy from (at least) three different spectroscopic surveys. This selection, coupled with our quality control requirement, essentially restricts our sample to sources whose calibration redshift is supported by multiple spectroscopic surveys with different selection functions, systematic effects, and catastrophic failure modes. This calibration sample is therefore expected to be very robust (albeit at some cost to statistical precision due to a significant reduction in photometric effective number density); mis-calibration of these data would require coordinated catastrophic failure of redshift assignment across multiple spectroscopic campaigns using different instruments and redshifting methods.

3. Results

The results of our various gold sample cosmic shear measurements, quantified using the marginal posterior constraints of the cosmic-shear summary parameter of interest $S_8 = \sigma_8 \sqrt{\Omega_m/0.3}$, are shown in Figure 1. Also shown are the results from H20 and Planck-Legacy (Planck Collaboration et al. 2018), for comparison. The left panel is split into three sections: analyses performed with more informative non-zero mean redshift bias priors (' $\delta z \neq 0$ ', see Appendix B), estimated using the simulations of van den Busch et al. (2020); analyses performed with the broader, zero-mean redshift bias priors of H20 (' $\delta z = 0$ '; used in cases where we cannot derive improved priors or wish to compare directly to previous work); and external results taken directly from the literature.

First, we verify our updated cosmology pipeline by performing an identical cosmological analysis to H20. As seen by the two results highlighted by the blue box in Figure 1, we find that we recover essentially the same S_8 as they report: $S_8 = 0.739^{+0.040}_{-0.037}$ (labelled 'KV450-DIR' in the figure, with $\delta z = 0$) compared to their $S_8 = 0.737^{+0.040}_{-0.036}$ ('Hildebrandt+ (2020)'). We argue that the observed difference ($|\Delta S_8| \lesssim 0.003$) is simply a reflection of noise within our Markov-chain Monte-Carlo (MCMC). For our fiducial gold sample analysis, shown in orange in both panels, we find a marginal constraint of $S_8 = 0.716^{+0.043}_{-0.038}$; smaller than, but otherwise consistent with, that which was found by H20. Our fiducial analysis is in better agreement with the results of H20 when their data set and redshift distributions are re-analysed with more informative redshift distribution bias priors (estimated from simulations, see Appendix B): $S_8 = 0.727^{+0.039}_{-0.036}$ ('KV450-DIR' with $\delta z \neq 0$; purple). We observe that our fiducial analysis has a slightly broader marginal S_8 constraint. This is expected when performing our gold selection: by decreasing the size of the photometric data set which is used for the analysis (which we quantify using the change in the effective number density of cosmic shear source galaxies, $\Delta n_{\text{eff}} = n_{\text{eff}}^{\text{gold}} / n_{\text{eff}}^{\text{all}} \approx 80\%$ for our fiducial sample; see Appendix C), we increase the statistical noise on our marginal constraints.

We explore the sensitivity of our analysis to the construction of our spectroscopic compilation, by performing our analysis with gold samples constructed without spec-

tra from zCOSMOS, VVDS, and DEEP2. When removing zCOSMOS or VVDS, we find that our marginal constraint on S_8 is unchanged within MCMC noise: $|\Delta S_8| \lesssim 0.003$. In the cases of removing DEEP2 from the calibration sample, we find a shift in our marginal constraint of $S_8 = 0.707^{+0.046}_{-0.042}$, equating to $|\Delta S_8| \lesssim 0.2\sigma$. We note though, that (looking at the Ω_m versus S_8 plane) we can see that the shift in S_8 without DEEP2 is driven by an extension of the posterior to lower values, rather than a systematic biasing of the distribution overall. This result is of further interest in the context of the full multi-dimensional posterior space, which we present in Appendix D for the interested reader.

We draw particular attention to the differences seen between our analysis without DEEP2 and the equivalent analysis performed by H20. When performing their noDEEP2 analysis H20 found a non-trivial increase in S_8 to $S_8 = 0.761^{+0.041}_{-0.037}$ ('H+20 noDEEP2'); a shift of $\Delta S_8 \sim +0.6\sigma$. This difference is attributed, in H20, to a bias in the reconstructed redshift distributions used for this test: removing DEEP2 causes pathological misrepresentation of the high-redshift portion of the spectroscopic colour-colour space, which subsequently causes the reconstructed redshift distributions to be systematically biased low, thereby introducing a positive shift in S_8 for the otherwise unchanged photometric source sample. To this end, H20 argued that their noDEEP2 analysis was systematically biased, and was a demonstration of the limitations of their redshift calibration procedure. In the SOM based gold samples, however, we find no such biasing of our results when calibrating without spectra from DEEP2. We attribute this difference to the gold selection process: our redshift distributions without DEEP2 are similarly skewed low compared to the fiducial (see Appendix C), however unlike H20 our gold selection simultaneously removes the misrepresented photometric sources. Therefore, while our redshift distributions change significantly between the fiducial and noDEEP2 analyses, both correctly describe the photometric data within their respective gold samples; both are accurate and consistent. We therefore see no biasing of the derived cosmological parameters, but rather just an increase in marginal uncertainties due to the aforementioned decrease in statistical power due to the $\sim 20\%$ reduction in the effective number density of the photometric sample.

In addition to our tests for the effect of pathological colour misrepresentation, we also test the influence of spectra which may have an increased fraction of catastrophic failures. Recall that in our spectroscopic compilation for KV450 we allow only high-confidence ($\geq 95\%$) and/or 'certain' ($\geq 99.5\%$ confidence) spectra; however even high-confidence spectra may have catastrophic failures. In W20 we demonstrated using simulations that expected fractions of catastrophic spectroscopic failures were unlikely to bias calibration of redshift distributions in KV450. Nonetheless, here we explore the influence of the lower-confidence spectroscopy on our conclusions. Our 'speczquality4' gold sample is calibrated using only certain confidence redshifts. The resulting marginal constraints of S_8 differ from our fiducial results by $|\Delta S_8| = 0.011$. We therefore conclude that the presence of lower-confidence spectra in our calibration data set does not introduce significant biases in our fiducial marginal constraints of S_8 .

Our speczquality4 result is of additional interest in the context of recent work presented by Hartley et al. (2020). For DES (i.e. using fewer photometric bands than used in

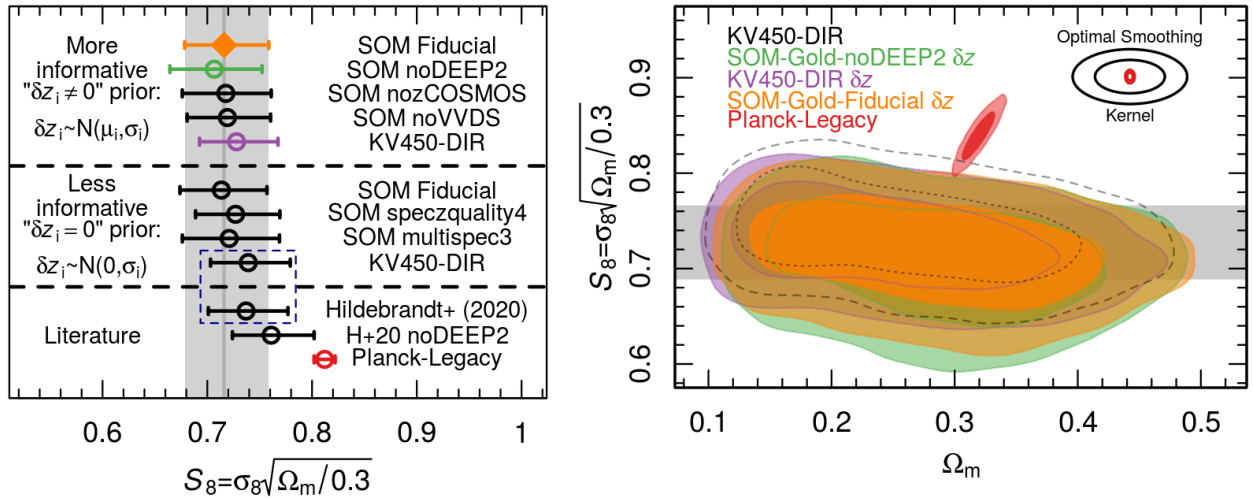


Fig. 1. Posterior constraints of S_8 (left) and Ω_m vs. S_8 (right) for our various gold samples, compared to the results from H20 and Planck CMB. We show results for our analyses split into sections, determined by the form of their redshift distribution priors. Results computed using more informative, non-zero mean, Gaussian redshift distribution bias priors ($\delta z \neq 0$, see Appendix B) for gold samples where these are able to be calculated, and using the zero mean Gaussian bias priors from H20 otherwise ($\delta z = 0$). For our fiducial analysis we show results with both priors, to allow direct comparison between our various results. We annotate our contour figure (right) with the two Gaussian smoothing kernels used in generating the contours (one for the cosmic shear contours, and one for the CMB contours). We find that our new cosmology pipeline produces results consistent with the pipeline of H20 (left panel, blue dashed box). Our fiducial results (orange) suggest a slightly lower S_8 than found in previous work: $S_8 = 0.716^{+0.043}_{-0.038}$. However we find that constraints on S_8 are extremely stable for all of our gold sample analyses here (compared to the fiducial: $|\Delta S_8| < 0.2\sigma$), demonstrating that the results here are more robust to spectroscopic misrepresentation than previous works. In particular, unlike H20, we find that even pathological misrepresentation at high-redshift (‘noDEEP2’) is unable to shift our estimates of S_8 to larger values.

KiDS), and implementing a redshift calibration methodology akin to that of H20, they find switching between direct calibration using high-confidence ($\geq 95\%$) and certain ($\geq 99.5\%$) spectroscopic samples results in a significant $\Delta(z) \geq 0.06$ bias for their highest tomographic bin ($Z_B \in (0.7, 1.3]$). While these biases are not directly applicable to our analysis, any similar systematic bias within our analysis would likely cause a significant change in the estimated cosmological parameters. We find no such significant bias when switching between direct calibration using high-confidence and certain spectroscopic redshifts, suggesting that this bias is suppressed in our data set. We hypothesise that this is driven by one, or a combination, of the following three effects. Firstly, that our 9-band photometric space is more resilient to spectroscopic selection biases than the 4-band space considered in Hartley et al. (2020). Secondly, that our deeper and more diverse spectroscopic compilation reduces the sensitivity of the recalibration procedure to strong (survey-specific) spectroscopic selection effects. Finally, that the calibration method of Wright et al. (2020) is more resilient to spectroscopic selection effects than the method used in Hartley et al. (2020). We leave exploration of these three possibilities to future work.

A change in marginal S_8 constraints of $|\Delta S_8| = 0.011$, with respect to the fiducial case, is nonetheless large in the context of our gold samples. As such, we extend this test further by implementing more stringent requirements on spectroscopic agreement. Our ‘multispec3’ gold sample consists only of photometric sources which are calibrated by spectra originating from at least 3 different spectroscopic surveys within our compilation. As stated in Section 2, this requirement places a strong restriction on spectroscopic agreement when coupled with our quality control re-

quirement (Equation 1). For our multispec3 gold sample we find again a result which is consistent with our fiducial analysis: $S_8 = 0.721^{+0.048}_{-0.045}$, corresponding to $|\Delta S_8| \leq 0.005$, only slightly larger than the MCMC noise threshold. This result demonstrates the stability of our gold sample results, and provides a strong indication that the marginal constraints on S_8 presented here are not biased by systematic effects nor catastrophic failures within the spectroscopic calibration sample.

While we have focussed our discussion here on the marginal S_8 constraints, in Appendix D we provide additional marginal and two-dimensional posterior constraints for a subset of our parameter distributions. We explore other conclusions which we can draw from our gold cosmological analyses, specifically around intrinsic alignments and the posterior probability distributions of Ω_m and σ_8 . Briefly, our gold sample marginal constraints show a reduced preference for low values of Ω_m , causing a more consistent recovery of $\Omega_m \approx 0.3$. In all of our gold analyses the marginal constraints are good agreement ($|\Delta X| < 0.2\sigma$ for all parameters X), with the exception of the intrinsic alignment amplitude parameter A_{IA} , which shows up to $|\Delta A_{IA}| \sim 1.0\sigma$ differences among analyses. Importantly, though, our gold sample A_{IA} constraints are all consistent with $A_{IA} = 0$, unlike those from H20, who found $A_{IA} \approx 1$. This updated constraint is in better agreement with recent work on intrinsic alignments (Fortuna et al. 2020), who predict an intrinsic alignment amplitude for KiDS of $0 \leq A_{IA} \leq 0.2$. Finally, we find that differences in S_8 and A_{IA} are correlated between our gold sample analyses, and that stronger prior information on A_{IA} would further improve the agreement between our respective gold sample analyses.

4. Summary

We present updated cosmological parameter constraints from the KiDS+VIKING-450 data set of Wright et al. (2019), estimated using updated redshift distributions following the methodology of Wright et al. (2020). For our fiducial analysis we find a value of S_8 that is smaller than, but nonetheless fully consistent with, the value reported in the previous KiDS+VIKING-450 cosmological analysis of Hildebrandt et al. (2020): $S_8 = 0.716^{+0.043}_{-0.038}$ compared to $S_8 = 0.737^{+0.040}_{-0.036}$ ($|\Delta S_8| \leq 0.6\sigma$). We note, however, that when one analyses the data and redshift distributions of Hildebrandt et al. (2020) using updated redshift distribution bias parameters presented in Wright et al. (2020), their S_8 also shifts downward and is in better agreement with our fiducial analysis: $S_8 = 0.727^{+0.039}_{-0.036}$, $|\Delta S_8| \leq 0.3\sigma$. We explore the sensitivity of our results to systematic misrepresentation within the spectroscopic calibration data set by removing multiple spectroscopic subsamples (DEEP2, VVDS, zCOSMOS), each of which uniquely calibrate distinct portions of the colour-redshift space. We find that the results presented here are robust to pathological misrepresentation, whereby even the removal of DEEP2 is unable to cause a significant shift in S_8 : $|\Delta S_8| \leq 0.2\sigma$. We find that our results are consistent with the fiducial when performing the calibration using only certain ($nQ = 4$, $\geq 99.5\%$ confidence) spectroscopic redshifts, and when performing an extremely conservative analysis considering only photometric sources which are simultaneously calibrated by spectra from at least three different spectroscopic surveys: $|\Delta S_8| \leq 0.011$ and $|\Delta S_8| \leq 0.005$ respectively. Overall these results indicate that, using the redshift calibration methodology of Wright et al. (2020), pathological misrepresentation of photometric sources within the spectroscopic compilation is not able to produce significant changes in marginal constraints of S_8 , and therefore cannot reconcile the $\Delta S_8 \approx 2.5\sigma$ differences observed between cosmological parameters estimated using KiDS and Planck.

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Appendix A: Cosmology and wrapper pipelines

With this letter we release a new implementation of the KiDS cosmological analysis pipeline utilised by H20, which has been generalised for ease of use. The new pipeline, simply called *CosmoPipe*, is available from <https://www.github.com/AngusWright/CosmoPipe>. The package can be installed trivially with the provided master installation script, and generates a clean working environment for each installation of the pipeline to avoid conflicts with, for example, existing python installations.

CosmoPipe contains the same analysis steps performed by H20. The pipeline utilises *treecorr* (Jarvis et al. 2004) for computation of cosmic shear correlation functions, and *MontePython* (Audren et al. 2013) for Markov-chain monte-carlo (MCMC) analyses. For clarity, we outline the seven primary steps of the pipeline here.

1. Compute the 2D c-term for all survey patches and tomographic bins;
2. Compute the 1D c-term for all survey patches;
3. Compute 2pt shear correlation functions;
4. Construct the correlation function covariance matrix;
5. Prepare the data for input to *MontePython* MCMC: reformat the correlation functions, reformat the covariance matrix, prepare the montepython likelihood, reformat the N_z distributions, define the correlation function scale-cuts, and link any required *treecorr* files;
6. Run the MCMC;
7. Construct summary figures and statistics from the MCMC chains.

While this pipeline has been largely generalised, it is clear that some of these steps above are tailored for KiDS-like cosmological analyses. For example, the *CosmoPipe* is provided with a version of the public KiDS likelihood that has been pre-formatted to fit seamlessly into the *CosmoPipe*. The code will function equally well with an arbitrary likelihood, albeit with some additional preparation required on the user-side.

Should one wish to perform an analysis such as (or indeed identical to) that presented here, we also provide a wrapper package which links together the cosmological analysis pipeline package and the redshift calibration package of W20. This wrapper package, available at <https://www.github.com/AngusWright/CosmoWrapper>, contains one main script, *Wright2020b.sh*, which performs the entirety of the analysis presented here. This script requires only that the user have the input photometric and spectroscopic calibration data sets supplied, and performs (with one command) the full gambit of analysis required for this letter. These steps include:

1. redshift calibration;
2. gold sample selection;
3. installation of *CosmoPipe*;
4. preparation of *CosmoPipe* for the different gold sample runs;
5. running *CosmoPipe*; and
6. outputting of figures present in this paper.

Some additional input parameters to the *CosmoPipe* are also encoded in the wrapper package, such as the various redshift distribution and multiplicative shear bias priors given discussed in Appendix B.

Appendix B: Gold sample priors

Appendix B.1: Mean redshift distribution biases

Redshift distribution bias is a significant uncertainty in cosmic shear analyses, as coherent bias in redshift distribution means is directly reflected as a bias in S_8 . In H20, redshift distributions were assigned a zero-mean ($\delta z = 0$) Gaussian prior, whose width (per tomographic bin) was estimated via a spatial bootstrap method (see Section 3.2 of H20). These priors are applicable to the direct calibration implemented in their work, but need not be reflective of the biases in our SOM based redshift calibration procedure.

Fortunately, W20 present estimates of the redshift calibration biases for our fiducial gold sample. Briefly, these redshift distribution bias estimates are made using simulations of the KiDS photometric and spectroscopic compilations from van den Busch et al. (2020). These simulations generate 100 independent realisations of the three main spectroscopic lines-of-sight within KiDS, which are then used to calibrate a mock KiDS photometric sample. At all steps in this process, samples are constructed to match the real KiDS photometric and spectroscopic data in all possible manners: magnitude, colour, selection effects, weights, redshift distributions, etc. This simulation therefore allows the estimation of mean redshift distribution biases, include uncertainties and biases introduced by photometric noise, sample variance, spectroscopic selection effects, spectroscopic incompleteness, and Poisson sampling. Tests performed in W20 suggest that the dominant sources of calibration uncertainty are photometric noise and spectroscopic selection effects. Using the recovered mean redshift distribution bias estimates (and uncertainties), we are able to construct new, more informative, non-zero-mean ($\delta z \neq 0$) redshift distribution bias priors for our cosmic shear analysis. These priors are presented in Table B.1.

As each of the zCOSMOS, VVDS, and DEEP2 samples were simulated by van den Busch et al. (2020), we are able to also calibrate the redshift bias priors for the three gold samples which exclude these subsamples: our ‘nozCOSMOS’, ‘noVVDS’, and ‘noDEEP2’ gold samples. However, the simulations of van den Busch et al. (2020) did not include estimates of redshift quality, so we are unable to construct a more informative prior for the ‘specquality4’ gold sample. Similarly, as they do not simulate every spectroscopic subsample within KiDS, construction of a new ‘multispec3’ prior is also not possible. In each of these cases, we are therefore required to default back to the $\delta z = 0$ priors of H20.

To be conservative, we opt to double the uncertainties on the bias found in the simulations when constructing our priors. We opt to utilise these new priors, where possible, as they represent our current best-estimate of the true redshift bias parameters inherent to the recalibration method and samples used here, despite the limitations of the simulations used (van den Busch et al. 2020; Wright et al. 2020). We note, however, that the biases are typically small, being of order $\delta z \lesssim 0.01$ for the majority of samples and bins, meaning that the H20 $\delta z = 0$ prior is nonetheless a reasonable approximation (and therefore unlikely to be a source of bias in the cases where we are required to use that prior; i.e. specquality4 and multispec3).

Table B.1. Updated redshift distribution bias priors parameters used in different gold sample analyses. Priors are Gaussian ($\mu \pm \sigma$). Parameters are determined from the simulations of [van den Busch et al. \(2020\)](#) as described in [W20](#), except that we double the simulation bias standard deviations when constructing our priors. For samples where we want to replicate previous analyses, we implement the prior from [H20](#) (‘All $\delta z = 0$ ’).

Gold Sample	Tomographic Redshift Bias Prior δz				
	bin1 $Z_B \in (0.1, 0.3]$	bin2 (0.3, 0.5]	bin3 (0.5, 0.7]	bin4 (0.7, 0.9]	bin5 (0.9, 1.2]
Fiducial	0.000 \pm 0.010	0.002 \pm 0.012	0.013 \pm 0.012	0.011 \pm 0.008	−0.006 \pm 0.010
KV450-DIR	0.047 \pm 0.010	0.025 \pm 0.008	0.032 \pm 0.010	−0.004 \pm 0.008	−0.013 \pm 0.008
$\delta z \neq 0$ NoDEEP2	−0.001 \pm 0.010	0.002 \pm 0.012	−0.002 \pm 0.012	−0.009 \pm 0.010	−0.015 \pm 0.010
noVVDS	0.001 \pm 0.010	0.001 \pm 0.012	0.024 \pm 0.014	0.014 \pm 0.010	−0.007 \pm 0.012
nozCOSMOS	0.005 \pm 0.026	0.005 \pm 0.016	0.032 \pm 0.014	0.030 \pm 0.010	0.002 \pm 0.012
All $\delta z = 0$	0.000 \pm 0.039	0.000 \pm 0.023	0.000 \pm 0.026	0.000 \pm 0.012	0.000 \pm 0.011

Appendix B.2: Multiplicative shear bias

As each of our gold selections produces a different subset of the full photometric sample, this requires a new computation of the multiplicative and additive shear biases, shear correlation functions, covariances, etc. Each of these is incorporated into the pipeline processing, with the exception of the multiplicative shear bias estimation. In [H20](#), multiplicative shear biases are computed using the methodology and simulations of [Kannawadi et al. \(2019\)](#). We invoke the same procedure, thereby generating a bespoke set of multiplicative shear-bias parameters for each of our gold selections, albeit outside of our wrapper pipeline. These bias parameters are given in Table B.2 for each of our gold samples. We note that in all cases we have chosen to implement the same m-bias uncertainty as used in [H20](#): $\Delta m = 0.02$ for all tomographic bins.

We recognise that this implementation of the m-bias estimation may be sub-optimal: the gold selections are strongly colour-dependent, and our current simulation-set for estimating multiplicative bias is entirely monochromatic (using only the *r*-band imaging and fluxes). However these simulations nonetheless represent the state-of-the-art within KV450, and we leave exploration of how the m-biases change with multi-colour simulations for future studies. These m-bias parameters are required as input to our cosmology pipeline, and so are documented here. Overall, the gold sample multiplicative biases are very similar, with only the multispec3 calibration differing from the fiducial by more than $\Delta m \sim 0.002$. In all cases, the different m-bias values are well within the assumed multiplicative bias uncertainty used here.

Appendix C: Gold sample representation statistics

In this work we have tested the sensitivity of our cosmological parameter estimates to differently constructed gold samples within KV450. Each of these gold samples produces a subset of the available photometric data, and results in a different set of tomographic redshift distributions. We present these representation statistics and corresponding redshift distributions means here in Table C.1.

The combinations of mean redshift and representation statistics tells an interesting story regarding which photometric data are being removed by each of our gold sample definitions. There is a clear correlation between the removal of photometric data and a subsequent decrease in the mean redshift of the tomographic bins. The most obvious examples of this are in the cases of our noDEEP2 and multispec3

samples, where the gold selection removes 30% and 45% of the fiducial n_{eff} in the fifth tomographic bin, respectively. These samples also show the largest redshift distribution shifts within our gold samples: $\Delta\langle z \rangle \sim 0.05$ in the fifth tomographic bin. This suggests that the gold-sample definition is preferentially removing truly high-redshift sources from the photometric sample, as expected. This is indicative of the robustness of the joint redshift distribution estimation and gold selection; unlike the case of the redshift calibration in [H20](#), each combination of gold-sample and redshift distribution presented here is compatible, and differences in sample mean redshifts are not indications of bias in the redshift calibration methodology. This is an important distinction between the different redshift distributions presented here and in [H20](#).

Finally, we note the impact that the reduced effective number density of each gold sample has on our posterior constraint of S_8 . The multispec3 subsample, for example, has roughly 50% of the photometric n_{eff} of the KV450-DIR sample per tomographic bin, but shows only a $\sim 35\%$ larger uncertainty on S_8 . This is in agreement with the results of [H20](#), who found that the KV450-DIR S_8 uncertainty was limited equally by statistical and systematic uncertainties.

Appendix D: Additional marginal constraints

Here we present a subset of the posterior constraints from a subset of our gold sample analyses. In Figures D.1 and D.2 we show four of the 14 cosmological and nuisance parameters which are used by our likelihood model (A_{IA} , n_s , h , and $\ln 10^{10} A_s$), as well as three derived parameters (Ω_m , σ_8 , and S_8), in their marginal and 2D projections respectively. The marginal mean and standard deviations of these posterior distributions are also provided in Table D.1. For an in depth description of the likelihood used here see [H20](#). We have selected these parameters to show as they are of cosmological interest and/or are not prior dominated in our analysis. In particular, we opt not to show the various redshift distribution bias posteriors, as they follow their priors exactly in all our gold sample analyses. In this way, our results reinforce those of previous works: cosmic shear correlation functions cannot self-calibrate redshift distribution biases.

The marginal distributions from each of our gold samples in Figure D.1 are in good agreement. Comparing the various gold sample analyses to our two ‘KV450-DIR’ runs, which use data vectors and redshift distributions equivalent to those in [H20](#), we see some interesting differences. Firstly, we note that the gold samples no longer demon-

Table B.2. Multiplicative shear bias parameters used for each of our gold sample analyses.

Gold Sample	multiplicative shear bias parameter				
	bin1	bin2	bin3	bin4	bin5
Fiducial	-0.0145 ± 0.0200	-0.0176 ± 0.0200	-0.0125 ± 0.0200	0.0045 ± 0.0200	0.0122 ± 0.0200
NoDEEP2	-0.0137 ± 0.0200	-0.0162 ± 0.0200	-0.0112 ± 0.0200	0.0054 ± 0.0200	0.0130 ± 0.0200
noVVDS	-0.0143 ± 0.0200	-0.0172 ± 0.0200	-0.0116 ± 0.0200	0.0047 ± 0.0200	0.0125 ± 0.0200
nozCOSMOS	-0.0143 ± 0.0200	-0.0159 ± 0.0200	-0.0106 ± 0.0200	0.0053 ± 0.0200	0.0135 ± 0.0200
speczquality4	-0.0141 ± 0.0200	-0.0163 ± 0.0200	-0.0121 ± 0.0200	0.0043 ± 0.0200	0.0125 ± 0.0200
multispec3	-0.0158 ± 0.0200	-0.0203 ± 0.0200	-0.0173 ± 0.0200	-0.0033 ± 0.0200	-0.0012 ± 0.0200

Table C.1. Mean tomographic redshifts and representation statistics of photometric source galaxies within each of our gold samples. Representation is defined using the effective number density of sources for cosmic shear studies, n_{eff} , in each of the gold samples relative to a reference sample $n_{\text{eff}}^{\text{fid}}$. For the fiducial representation statistic we use the full KV450 photometric data set for reference (i.e. $n_{\text{eff}}^{\text{fid}}/n_{\text{eff}}^{\text{all}}$), while all other gold sample representations use the fiducial for reference (i.e. $n_{\text{eff}}^{\text{gold}}/n_{\text{eff}}^{\text{fid}}$). The statistics are all given per tomographic bin. The table demonstrates that each of our nozCOSMOS, noVVDS, and noDEEP2 gold samples has preferentially removed a different section of the colour-space. This is joined, however, by a shift in the mean redshift of the tomographic bin, indicating that the loss of the colour redshift space has been accounted for in the reconstruction. As expected, the multispec3 selection is highly restrictive, removing 30 – 45% of the fiducial photometric n_{eff} in every bin.

Gold Sample	$n_{\text{eff}}^{\text{gold}}/n_{\text{eff}}^{\text{ref}}(\%)$					$\langle z \rangle$				
	bin1	bin2	bin3	bin4	bin5	bin1	bin2	bin3	bin4	bin5
KV450-DIR	100.0	100.0	100.0	100.0	100.0	0.369	0.463	0.643	0.806	0.973
Fiducial	78.6	82.1	79.2	82.3	91.6	0.236	0.379	0.537	0.766	0.948
nozCOSMOS	93.4	92.2	92.0	88.3	91.5	0.214	0.371	0.529	0.755	0.945
noDEEP2	97.7	96.2	88.6	79.5	72.7	0.237	0.374	0.516	0.737	0.908
noVVDS	97.1	92.4	86.2	88.1	91.4	0.237	0.373	0.537	0.766	0.951
speczquality4	95.0	92.0	87.2	86.8	89.4	0.231	0.367	0.524	0.756	0.941
multispec3	71.2	72.7	65.0	55.4	54.5	0.226	0.369	0.515	0.737	0.906

Table D.1. Marginal parameter means and standard deviations for the subset of parameters shown in Figure D.1.

Parameter	KV450-DIR	Fiducial	nozCOSMOS	noVVDS	noDEEP2	KV450-DIR
			$\delta z \neq 0$			$\delta z = 0$
A_{IA}	0.282 ± 0.594	-0.344 ± 0.695	-0.366 ± 0.650	0.198 ± 0.665	-0.627 ± 0.775	0.959 ± 0.671
n_s	1.044 ± 0.130	1.023 ± 0.133	1.042 ± 0.136	1.020 ± 0.137	1.072 ± 0.130	1.032 ± 0.131
h	0.747 ± 0.049	0.742 ± 0.049	0.743 ± 0.049	0.741 ± 0.050	0.741 ± 0.049	0.748 ± 0.048
$\ln 10^{10} A_s$	3.158 ± 0.864	2.816 ± 0.806	2.762 ± 0.795	2.617 ± 0.715	2.653 ± 0.743	3.099 ± 0.882
Ω_m	0.249 ± 0.082	0.282 ± 0.085	0.286 ± 0.085	0.305 ± 0.085	0.291 ± 0.081	0.259 ± 0.087
σ_8	0.834 ± 0.156	0.768 ± 0.143	0.765 ± 0.145	0.739 ± 0.132	0.743 ± 0.134	0.833 ± 0.160
S_8	0.728 ± 0.035	0.716 ± 0.038	0.718 ± 0.041	0.719 ± 0.038	0.707 ± 0.042	0.739 ± 0.036

strate a preference for small values of the matter density parameter, $\Omega_m \sim 0.18$. Instead, our gold marginal distributions all peak at values $\Omega_m \sim 0.3$, in much better agreement with concordance cosmological parameters. This has a subsequent effect on the marginal distribution of σ_8 , causing it to be considerably narrower for our gold analysis than in the KV450-DIR cases; we find $\sigma_8 = 0.762^{+0.070}_{-0.180}$ compared to $\sigma_8 = 0.836^{+0.132}_{-0.218}$.

Focussing on the marginal constraints on A_{IA} , we see that this parameter shows the greatest variation within our gold sample analyses. Interestingly, though, we note that only the results of ‘KV450-DIR’ (i.e. KV450-DIR using the fiducial redshift bias priors) demonstrate a preference for non-zero values of A_{IA} . In all other cases, the marginal constraints are consistent with $A_{\text{IA}} = 0$; in agreement with recent work on intrinsic alignments within KiDS (Fortuna et al. 2020), who advocate $0 \leq A_{\text{IA}} \leq 0.2$.

Furthermore, we note a correlation between the values of A_{IA} and S_8 in the marginal distributions. Looking to the 2D posteriors in Figure D.2, the degeneracy between these two parameters is clearer. Interestingly, the A_{IA} versus S_8

joint posterior (for the ‘KV450-DIR’, ‘SOM-Gold Fiducial’, and ‘SOM-Gold noDEEP2’ samples shown) demonstrates that a stronger prior on the value of A_{IA} would likely bring the value of S_8 in these analyses into even closer agreement. This in turn suggests that the freedom of our intrinsic alignment parameter does not improve the agreement of S_8 between our different gold samples; rather the opposite. For demonstrative purposes only, we present in Table D.2 the marginal posterior constraints on S_8 for posterior samples within $0 \leq A_{\text{IA}} \leq 0.2$, compared to the full posterior results (which invokes an intrinsic alignment amplitude prior of $-6 \leq A_{\text{IA}} \leq 6$). This clearly demonstrates the interplay between A_{IA} and S_8 ; for the restricted prior, we find a maximal deviation between marginal constraints of $|\Delta S_8| \lesssim 0.005$, only marginally above MCMC noise. Nonetheless, we caution against the over-interpretation of these constraints, as a stronger prior on A_{IA} is not yet justifiable: analyses such as Fortuna et al. (2020) invoke a number of assumptions that prohibit the use of their constraints as prior distributions.

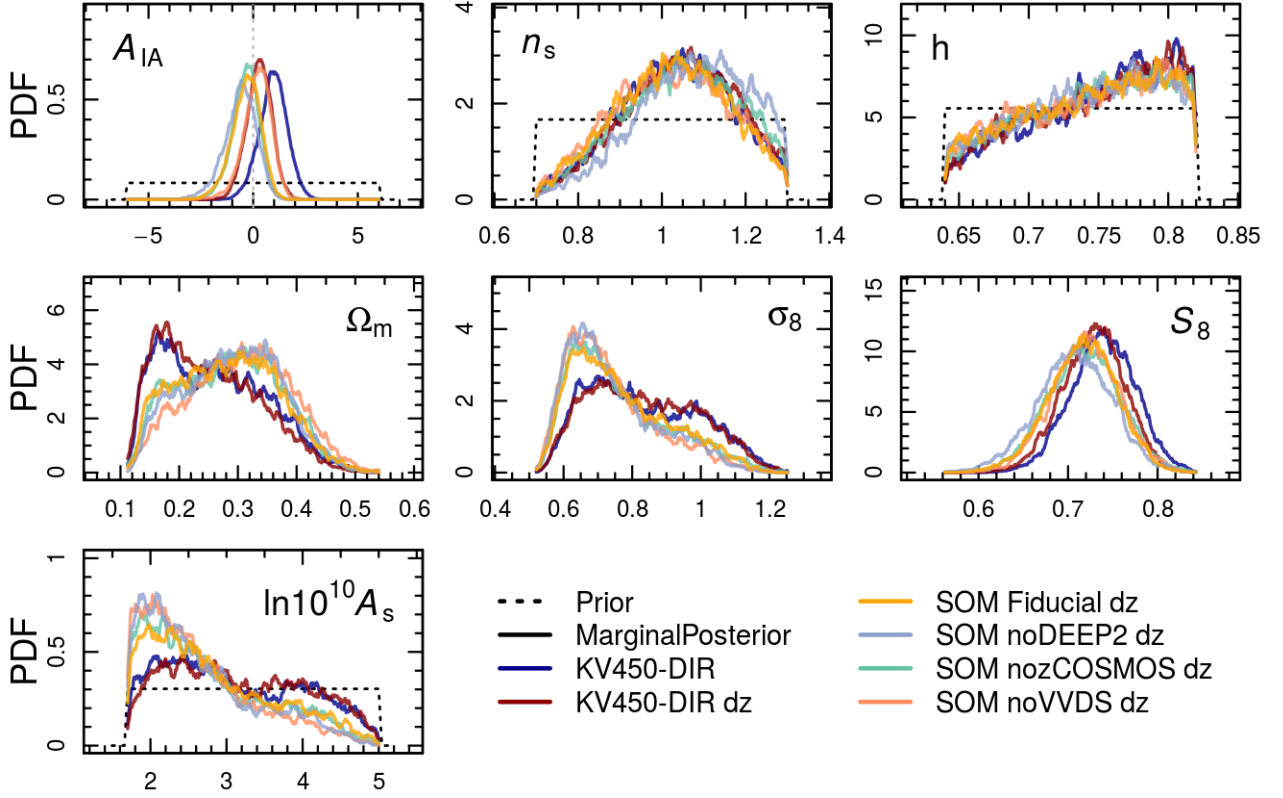


Fig. D.1. Marginal posterior distributions for a subset of the cosmological, nuisance, and derived parameters used in our cosmological model. Coloured lines represent the marginal distributions from various samples. Dashed lines show the priors for all non-derived parameters. There is a clear difference between the marginal distributions of the gold and full-sample (‘KV450-DIR’) analyses. We note in particular that the previously observed preference within KV450 for small values of the matter density parameter Ω_m is removed in our gold analyses. The gold analyses also prefer a lower value of A_{IA} , consistent with 0 in all cases.

Table D.2. For demonstrative purposes only, the marginal constraints of S_8 with our fiducial and reduced A_{IA} prior. This demonstrates that, at fixed values of A_{IA} , each of the gold samples are in agreement with respect to S_8 , and therefore that the differences seen in S_8 are not improved by the freedom of A_{IA} in our cosmological model.

	Gold Sample	Marginal S_8 Constraint	
		$-6 \leq A_{IA} \leq 6$	$0 \leq A_{IA} \leq 0.2$
$\delta z \neq 0$	KV450-DIR	0.728 ± 0.035	0.729 ± 0.035
	Fiducial	0.716 ± 0.038	0.724 ± 0.036
	nozCOSMOS	0.718 ± 0.041	0.728 ± 0.038
	noVVDS	0.719 ± 0.038	0.722 ± 0.037
	noDEEP2	0.707 ± 0.042	0.724 ± 0.038
$\delta z = 0$	KV450-DIR	0.739 ± 0.036	0.730 ± 0.034
	Fiducial	0.714 ± 0.039	0.722 ± 0.036
	speczquality4	0.727 ± 0.039	0.729 ± 0.037
	multispec3	0.721 ± 0.044	0.725 ± 0.041

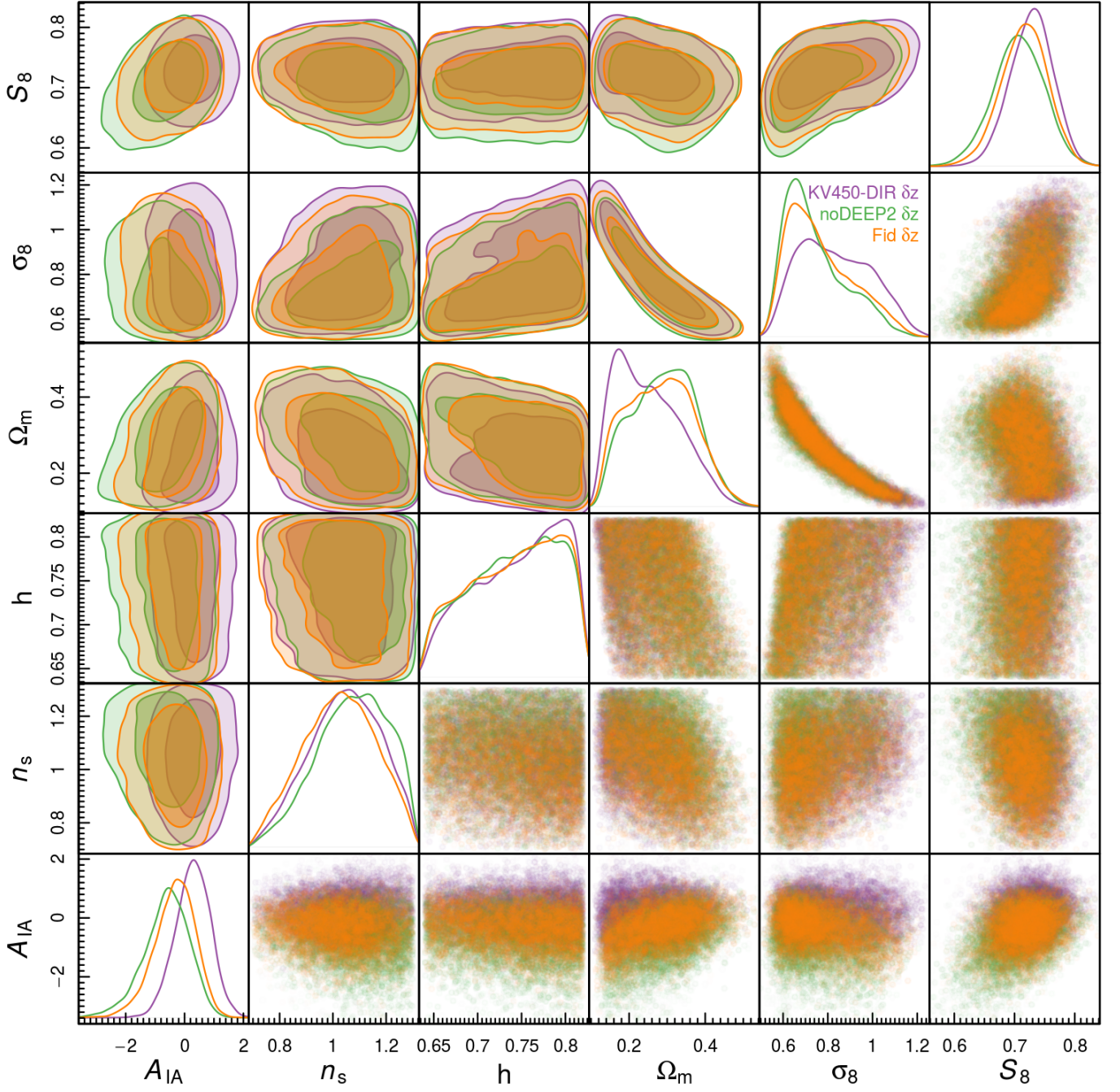


Fig. D.2. Marginal and 2D posterior distributions for a subset of the cosmological, nuisance, and derived parameters used in our cosmological model. We restrict this figure to only the ‘KV450-DIR δz ’ (purple), ‘SOM-Gold Fiducial δz ’ (gold), and ‘SOM-Gold noDEEP2 δz ’ (green) analyses, as these three analyses encompass the range of posteriors in all of our gold sample analyses (and show similar degeneracies). The joint posterior of A_{IA} and S_8 demonstrates that, were a stronger prior on A_{IA} justifiable, agreement between the various gold samples would be even higher.